

## Proceedings of the International Astronomical Union

Date of delivery: 5 May 2016

Journal and vol/article ref: IAU 1600006

Number of pages (not including this page): 6

This proof is sent to you on behalf of Cambridge University Press. Please check the proofs carefully. Make any corrections necessary on a hardcopy and answer queries on each page of the proofs

Please return the **marked proof** within **5** days of receipt to:

Managing editor of this symposium

**Authors are strongly advised to read these proofs thoroughly because any errors missed may appear in the final published paper. This will be your ONLY chance to correct your proof. Once published, either online or in print, no further changes can be made.**

To avoid delay from overseas, please send the proof by airmail or courier.

If you have **no corrections** to make, please email **managing editor** to save having to return your paper proof. If corrections are light, you can also send them by email, quoting both page and line number.

- The proof is sent to you for correction of typographical errors only. Revision of the substance of the text is not permitted, unless discussed with the editor of the journal. Only **one** set of corrections are permitted.
- Please answer carefully any author queries.
- Corrections which do NOT follow journal style will not be accepted.
- A new copy of a figure must be provided if correction of anything other than a typographical error introduced by the typesetter is required.

**If you do not send any corrections to the editor within 5 days, we will assume your proof is acceptable.**

- If you have problems with the file please contact

**[lwebb@cambridge.org](mailto:lwebb@cambridge.org)**

Please note that this pdf is for proof checking purposes only. It should not be distributed to third parties and may not represent the final published version.

**Important:** you must return any forms included with your proof. We cannot publish your article if you have not returned your signed copyright form.

NOTE - for further information about **Journals Production** please consult our **FAQs** at [http://journals.cambridge.org/production\\_faqs](http://journals.cambridge.org/production_faqs)

---

**Author queries:**

---

**Typesetter queries:**

---

**Non-printed material:**

# Statistical properties of superflares on solar-type stars based on the Kepler 1-min cadence data

Hiroyuki Maehara<sup>1,\*</sup>, Takuya Shibayama<sup>2</sup>, Yuta Notsu<sup>3</sup>, Shota Notsu<sup>3</sup>, Satoshi Honda<sup>4</sup>, Daisaku Nogami<sup>3</sup> and Kazunari Shibata<sup>5</sup>

\*email: h.maehara@oao.nao.ac.jp

<sup>1</sup>Okayama Astrophysical Observatory, National Astronomical Observatory of Japan, 3037-5 Honjo, Kamogata, Asakuchi, Okayama, Japan, 719-0232

<sup>2</sup>Institute for Space-Earth Environmental Research, Nagoya University, Furo-cho, Chikusa-ku, Nagoya, Aichi, Japan, 464-8601

<sup>3</sup>Department of Astronomy, Kyoto University, Kitashirakawa-Oiwake-cho, Sakyo-ku, Kyoto, Japan, 606-8502

<sup>4</sup>Center for Astronomy, University of Hyogo, 407-2, Nishigaichi, Sayo-cho, Sayo, Hyogo, Japan, 679-5313

<sup>5</sup>Kwasan and Hida Observatories, Kyoto University, Yamashina-ku, Kyoto, Japan, 607-8471

**Abstract.** We searched for superflares on solar-type stars using the Kepler short-cadence (1-min sampling) data in order to detect superflares with short duration. We found 187 superflares on 23 solar-type stars whose bolometric energy ranges from the order of  $10^{32}$  erg to  $10^{36}$  erg. Using these new data combined with the results from the data with 30-min sampling, we found the occurrence frequency ( $dN/dE$ ) of superflares as a function of flare energy ( $E$ ) shows the power-law distribution ( $dN/dE \propto E^{-\alpha}$ ) with  $\alpha = 1.5$  for  $10^{33} < E < 10^{36}$  erg. The upper limit of energy released by superflares is basically comparable to a fraction of the magnetic energy stored near starspots which is estimated from the amplitude of brightness variations. We also found that the duration of superflares ( $\tau$ ) increases with the flare energy ( $E$ ) as  $\tau \propto E^{0.39 \pm 0.03}$ . This can be explained if we assume the time-scale of flares is determined by the Alfvén time.

**Keywords.** stars:activity, stars:flare, stars:solar-type

## 1. Introduction

Solar-flares are energetic eruption in the solar atmosphere caused by the magnetic reconnection (e.g. Shibata & Magara 2011). The occurrence frequency of solar-flares ( $dN/dE$ ) as a function of flare energy ( $E$ ) can be fitted by a simple power-law function ( $dN/dE \propto E^{-\alpha}$ ) with  $\alpha = 1.5 - 1.9$  in the flare energy range between  $10^{24}$  erg and  $10^{32}$  erg (e.g. Crosby *et al.* 1993, Shimizu 1995, Aschwanden *et al.* 2000). The bolometric energy released by an X10 class solar flare is estimated to be the order of  $10^{32}$  erg (Emslie *et al.* 2012) and the occurrence frequency of such flare is about once in 10 years. In the case of other stars, including solar-type stars, more energetic flares called “superflares” have been observed (e.g. Landini *et al.* 1986, Schaefer 1989, Schaefer *et al.* 2000). Recently, many superflares on solar-type stars (G-type main sequence stars) have been discovered by the Kepler space telescope (e.g. Maehara *et al.* 2012, Shibayama *et al.* 2013). The bolometric energy released by the superflare ranges from  $10^{33}$  to  $10^{36}$  erg which is  $10\text{--}10^4$  times larger than that of the largest solar flares ( $\sim 10^{32}$  erg). Majority of solar-type stars exhibiting superflares show quasi-periodic brightness variations with amplitude range from  $\sim 0.1\%$  to  $8\%$ . These light variations are thought to be caused by the rotation of the star with starspots (e.g. Notsu *et al.* 2013). Although most of solar-type stars with superflares are

44 rapidly rotating stars, the temperature and rotation velocity of some superflare stars are  
 45 close to those of the Sun (Nogami *et al.* 2014).

46 Previous studies of superflares on solar-type stars using the Kepler long-cadence data  
 47 mainly focus on superflares with extremely large energy ( $E > 10^{34}$  erg) and long duration  
 48 ( $> 1$  hour). Here we report the results from the data with high time-resolution ( $\sim 1$   
 49 min) and discuss the statistical properties (e.g. occurrence frequency, flare duration) of  
 50 superflares on solar-type stars. More details are described in Maehara *et al.* (2015).

## 51 2. Data

52 We searched for flares from the short-cadence data ( $\sim 1$  min interval) observed with  
 53 the Kepler space telescope between 2009 April (quarter 0: Q0) and 2013 May (Q17)  
 54 (Koch *et al.* 2010, Gilliland *et al.* 2010). We selected solar-type (G-type main sequence)  
 55 stars from the data set by using the surface temperature of the star ( $T_{\text{eff}}$ ) and the surface  
 56 gravity ( $\log g$ ) taken from Huber *et al.* (2014) instead of those from Brown *et al.* (2011)  
 57 (initial Kepler Input Catalog). In previous papers (Maehara *et al.* 2012, Shibayama *et al.*  
 58 2013), we used stellar parameters taken from Brown *et al.* (2011) and the selection criteria  
 59 of  $5100\text{K} < T_{\text{eff}} < 6000$  K and  $\log g > 4.0$ . However the temperatures in Brown *et al.*  
 60 (2011) are systematically lower by  $\sim 200$  K than those in Huber *et al.* (2014). In order  
 61 to reduce the difference in statistical properties of superflares caused by the systematic  
 62 difference in stellar temperature, here we used the selection criteria of  $5300\text{K} < T_{\text{eff}} <$   
 63  $6300\text{K}$  and  $4.0 < \log g < 4.8$ . The total number of solar-type stars observed with short-  
 64 cadence mode is 1547.

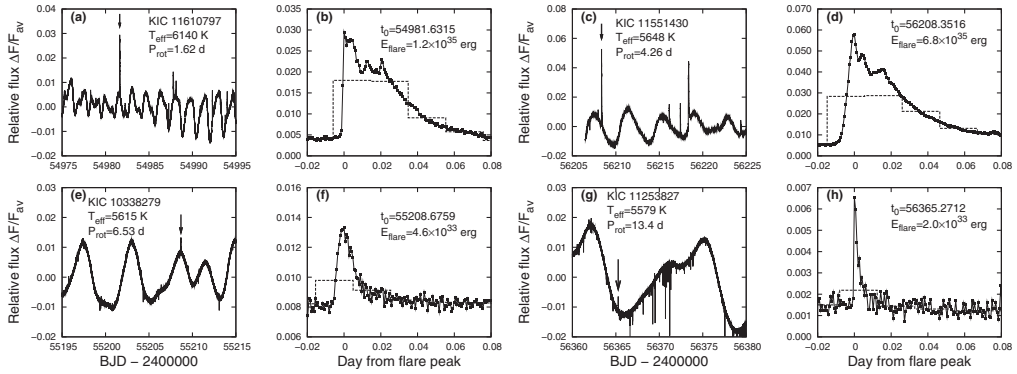
65 We used the Presearch Data Conditioned (PDC) light curve (Stumpe *et al.* 2012, Smith  
 66 *et al.* 2012) for the detection of flares. The detail procedures of the flare detection are  
 67 described in Maehara *et al.* (2015). Bolometric energy released by each flare was estimated  
 68 from the flare amplitude, flare duration, and stellar luminosity with the same manner  
 69 as Shibayama *et al.* (2013). The stellar luminosity were calculated from the effective  
 70 temperature and the stellar radius taken from Huber *et al.* (2014). We estimated periods  
 71 of long-term light variations from light curves of each star by using the discrete Fourier  
 72 transform (DFT) method. For period analysis, we used the long-cadence data (time  
 73 resolution of 30 min) obtained from 2009 September (Q2) to 2013 April (Q16). These  
 74 periods are basically consistent with those in McQuillan *et al.* (2014).

## 75 3. Results and discussion

76 We detected 187 flares on 23 solar-type stars from the data of 1547 solar-type stars.  
 77 Figure 1 shows typical light curves of superflares detected from short-cadence data. The  
 78 amplitude of flares normalized by the average brightness of the star range from  $1.3 \times 10^{-3}$   
 79 to  $8.5 \times 10^{-2}$  and the bolometric energies of flares range from  $2 \times 10^{32}$  to  $8 \times 10^{35}$  erg.

### 80 3.1. Occurrence frequency distribution

81 Figure 2 (a) represents the occurrence frequency of superflares ( $dN/dE$ ) as a function of  
 82 the bolometric energy of superflares ( $E$ ). Solid histogram in figure 2 (a) represents the  
 83 frequency distribution of superflares on all solar-type stars derived from short-cadence  
 84 data and dashed histogram represents that from long-cadence data (Shibayama *et al.*  
 85 2013). Since the period distribution of the stars observed in short-cadence mode is biased  
 86 to the shorter-period end, we estimate the corrected occurrence frequency for short-  
 87 cadence data taking into account the bias in the period distribution of the observed  
 88 samples. Both frequency distributions from short- and long-cadence data are almost the



**Figure 1.** (a), (c), (e), (g): Long-term light variations of the G-dwarf KIC 11610797 (a), KIC 11551430 (c), KIC 10338279 (e), and KIC 11253827 (g). The vertical axis means the relative difference between observed brightness of the star and the average brightness during the observation period. The horizontal axis means the times of the observations in Barycentric Julian Date. (b), (d), (f), (h): Enlarged light curve of a superflare on KIC 11610797 (b), KIC 11551430 (d), KIC 10338279 (f), and KIC 11253827 (h) (indicated by the down arrow in panel (a), (c), (e), and (g)). Filled-squares with solid-lines and dashed-lines represent the light curves from short- and long-cadence data respectively.

same for the flare energy between  $10^{34}$  erg and  $10^{36}$  erg and can be fitted by a power-law function ( $dN/dE \propto E^{-\alpha}$ ). Using the combined data set from both short- and long-cadence data, the power-law index  $\alpha$  is  $1.5 \pm 0.1$  for the flare energy of  $4 \times 10^{33} - 1 \times 10^{36}$  erg.

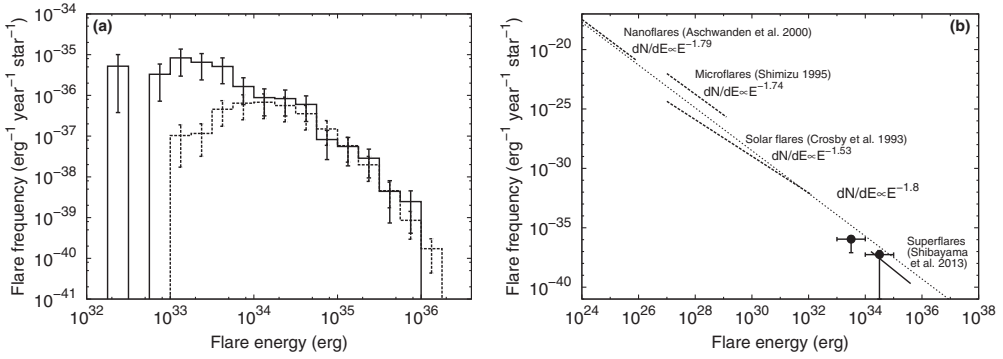
According to Shibata *et al.* (2013), the frequency distribution of superflares on Sun-like stars (early-G dwarfs with  $P_{\text{rot}} \geq 10$  day) and those of solar flares are roughly on the same power-law line. Figure 2 (b) represents the comparison between the frequency distribution of superflares on Sun-like stars ( $5800 \leq T_{\text{eff}} < 6300$  and  $P_{\text{rot}} \geq 10$  day) derived from both short- (filled circles) and long-cadence (solid line) data and those of solar flares (dashed lines). The thin dotted line indicates the power-law function ( $dN/dE \propto E^{-\alpha}$ ) with an index  $\alpha = 1.8$  taken from Shibata *et al.* (2013).

### 3.2. Flare energy and area of starspots

Majority of solar-type stars exhibiting superflares show quasi-periodic light variations with the amplitude of the order of 1 %. These variations are thought to be caused by the rotation of star with spotted surface and the amplitude of light variations corresponds to the area of starspots. Figure 3 shows the scatter plot of the energy of solar flares and superflares as a function of the area of sunspot/starspot group ( $A_{\text{spot}}$ ). The area of starspots on superflare stars were estimated from the amplitude of light variations. Since the flares are a result of a sudden conversion of magnetic energy to thermal and kinetic energy, the bolometric energy released by the flare should be a fraction of the magnetic energy ( $E_{\text{mag}}$ ) ((Shibata *et al.* 2013, Notsu *et al.* 2013)). The energy released by a flare can be written as the following form:

$$E_{\text{flare}} \sim f E_{\text{mag}} \sim \frac{f B^2 L^3}{8\pi} \sim \frac{f B^2 A_{\text{spot}}^{3/2}}{8\pi}, \quad (3.1)$$

where  $f$ ,  $B$  and  $L$  correspond to the fraction of magnetic energy released by the flare, the magnetic field strength of the starspots and the scale length of the starspot group respectively. The equation (3.1) suggests that the flare energy is roughly proportional to the area of the starspot group to the power of 3/2. According to Aschwanden *et al.* (2014),



**Figure 2.** (a): Bold solid and dashed histograms represent the occurrence frequency of superflares on all solar-type stars from short- (this work) and long-cadence data (Shibayama *et al.* 2013) as a function of total energy of superflares respectively. The vertical axis indicates the number of superflares per star, per year, and per unit energy. (b): Comparison between occurrence frequency superflares on Sun-like stars and those of solar flares. Filled-circles indicates the occurrence frequency of superflares on Sun-like stars (solar-type stars with  $P_{\text{rot}} > 10$  days and  $5800 < T_{\text{eff}} < 6300\text{K}$ ) derived from short-cadence data. Bold-solid line represents the power-law frequency distribution of superflares on Sun-like stars taken from Shibayama *et al.* (2013). Dashed lines indicate the power-law frequency distribution of solar flares observed in hard X-ray (Crosby *et al.* 1993), soft X-ray microflares (Shimizu 1995), and EUV nanoflaers (Aschwanden *et al.* 2000). Occurrence frequency distributions of superflares on Sun-like stars and solar flares are roughly on the same power-law line with an index of  $-1.8$  (thin-solid line; Shibata *et al.* 2013) for the wide energy range between  $10^{24}$  erg and  $10^{35}$  erg.

$f \sim 0.1$ . The typical sunspot area for generating X10-class flares ( $E_{\text{flare}} \sim 10^{32}$  erg) observed in 1989-1997 was  $3 \times 10^{-4}$  of the half area of the solar surface (Sammis *et al.* 2000), and the typical magnetic field strength of sunspot is the order of 1000 G (e.g. Antia *et al.* 2003). Therefore equation (3.1) can be written as

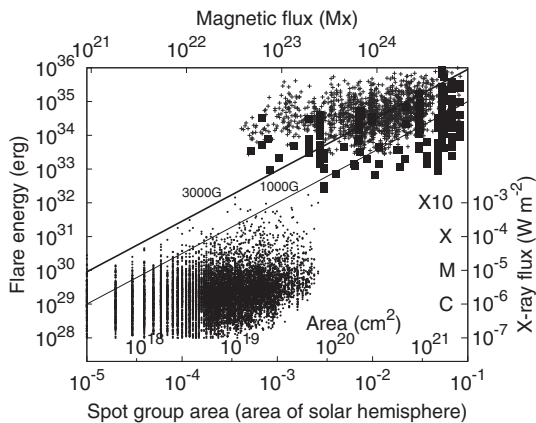
$$E_{\text{flare}} \sim 7 \times 10^{32}(\text{erg}) \left( \frac{f}{0.1} \right) \left( \frac{B}{1000\text{G}} \right)^2 \left[ \frac{A_{\text{spot}}/2\pi R_{\odot}^2}{0.001} \right]^{3/2}. \quad (3.2)$$

Solid and dashed lines in the figure 3 represent equation 3.2 for  $f = 0.1$ ,  $B = 1000$  G and  $f = 0.1$ , 3000 G respectively. Majority of superflares detected from short-cadence data (filled-squares) and almost all solar flares (small dots) are below these analytic lines. This suggests that the upper limit of the energy released by the flare is basically comparable to the stored magnetic energy estimated from the area of starspots. However there are some superflares above the analytic relation. These superflares would occur on the stars with low-inclination angle or stars with starspots around polar region. In the case of such stars, the light variation caused by the rotation become small. Hence the the actual area of starspots should be larger than that estimated from the amplitude of brightness variation. According to Notsu *et al.* (2015), some superflare stars which show superflares with energy larger than that expected from the amplitude of light variation were confirmed to have low-inclination angles. Moreover polar spots are often observed on rapidly-rotating late-type stars (e.g. Strassmeier 2009).

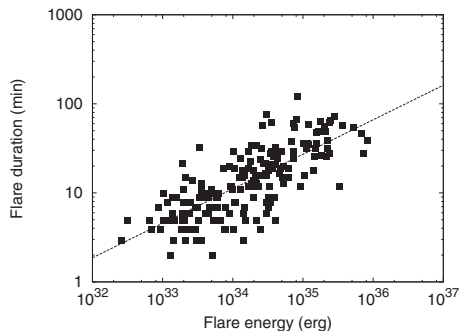
### 3.3. Correlation between flare duration and flare energy

Figure 4 represents the duration ( $e$ -folding decay time) of superflares as a function of the energy of flares. A linear fit for the superflares from short-cadence data in the log-log plot yields

$$\tau_{\text{flare}} \propto E_{\text{flare}}^{0.39 \pm 0.03} \quad (3.3)$$



**Figure 3.** Scatter plot of flare energy as a function of spot area. The lower and upper horizontal axis indicate the area of starspot group in the unit of the area of solar hemisphere and the magnetic flux for  $B = 3000$  G. The vertical axis represents the bolometric energy released by each flare. Filled-squares and small-crosses indicate superflares on solar-type stars from short-(this work) and long-cadence data (Shibayama *et al.* 2013) respectively. Small filled-circles represent solar flares (Ishii *et al.*, private communication, based on the data retrieved from the website of the National Geophysical Data Center of the National Oceanic and Atmospheric Administration (NOAA/NGDC), Solar-Terrestrial Physics Division at <http://www.ngdc.noaa.gov/stp/>).



**Figure 4.** Scatter plot of the duration of superflares as a function of the bolometric energy. Filled-squares indicate superflares on solar-type stars from short-cadence data. We used  $e$ -folding decay time as the flare duration. Dotted line indicates the linear regression for the data of superflares from short-cadence data. The power-law slope of the line is  $0.39 \pm 0.03$

where  $\tau_{\text{flare}}$  and  $E_{\text{flare}}$  indicate the duration and energy of superflares. Similar correlation between the flare duration and energy was observed in solar flares (e.g. Veronig *et al.* 2002, Christe *et al.* 2008). The flare energy is related to the magnetic energy stored near the starspots as shown in equation 3.1. On the other hand, the duration of impulsive phase of flares is comparable to the reconnection time ( $\tau_{\text{rec}}$ ) which can be written as

$$\tau_{\text{flare}} \sim \tau_{\text{rec}} \sim \tau_A / M_A \sim L / v_A / M_A, \quad (3.4)$$

where  $\tau_A = L / v_A$  is the Alfvén time,  $v_A$  is the Alfvén velocity, and  $M_A$  is the non-dimensional reconnection rate ( $M_A \sim 0.1$ - $0.01$  for the fast reconnection; Shibata & Magara 2011). If we assume  $B$  and  $v_A$  are not so different among solar-type stars, the duration of flares can be written as

$$\tau_{\text{flare}} \propto E_{\text{flare}}^{1/3}. \quad (3.5)$$

This suggests that the power-law slope for the correlation between the flare duration and flare energy is about  $1/3$  which is comparable to the observed value of  $0.39 \pm 0.03$ .

#### 4. Conclusions

We found 187 superflares on 23 solar-type stars from the Kepler 1-min cadence data. This new data set of superflares on solar-type stars suggests that the power-law frequency distribution of superflares (e.g. Maehara *et al.* 2012, Shibayama *et al.* 2013) continues to the flare energy of  $10^{33}$  erg. We also found the flare duration increases with the flare

energy as  $\tau_{\text{flare}} \propto E_{\text{flare}}^{0.39 \pm 0.03}$ . This suggests that the time-scale of flares is determined by the Alfvén time.

## References

- Antia, H. M., Bhatnagar, A., & Ulmschneider, P. 2003, Lectures on Solar Physics (Berlin: Springer Verlag) p. 133
- Aschwanden, M. J., Tarbell, T. D., & Nightingale, R. W., *et al.* 2000, *ApJ*, 535, 1047
- Aschwanden, M. J., Xu, Y., & Jing, J. 2014, *ApJ*, 797, 50
- Brown, T. M., Latham, D. W., Everett, M. E., & Esquerdo, G. A. 2011, *ApJ*, 142, 112
- Christe, S., Hannah, I. G., & Krucker, S., *et al.* 2008, *ApJ*, 677, 1385
- Crosby, N. B., Aschwanden, M. J., & Dennis, B. R. 1993, *Solar Phys.*, 143, 275
- Emslie, A. G., Dennis, B. R., & Shih, A. Y., *et al.* 2012, *ApJ*, 759, 71
- Gilliland, R. L., Jenkins, J. M., & Borucki, W. J., *et al.* 2010, *ApJ* (Letters), 713, L160
- Huber, D., Silva Aguirre, V., & Matthews, J. M., *et al.* 2014, *ApJS*, 211, 2
- Koch, D. G., Borucki, W. J., & Basri, G., *et al.* 2010, *ApJ* (Letters) 713, L79
- Landini, M., Monsignori Fossi, B. C., Pallavicini, R., & Piro, L. 1986, *A&A*, 157, 217
- Maehara, H., Shibayama, T., & Notsu, S., *et al.* 2012, *Nature*, 485, 478
- Maehara, H., Shibayama, T., & Notsu, Y., *et al.* 2015, *Earth, Planet and Space*, 67, 59
- McQuillan, A., Mazeh, T., & Aigrain, S. 2014, *ApJS*, 211, 24
- Nogami, D., Notsu, Y., & Honda, S., *et al.* 2014, *PASJ*, 66, L4
- Notsu, Y., Shibayama, T., Maehara, H., *et al.* 2013, *ApJ*, 771, 127
- Notsu, Y., Honda, S., & Maehara, H., *et al.* 2015, *PASJ*, 67, 33
- Sammis, I., Tang, F., & Zirin, H. 2000, *ApJ*, 540, 583
- Schaefer, B. E. 1989, *ApJ*, 337, 927
- Schaefer, B. E., King, J. R., & Deliyannis, C. P. 2000, *ApJ*, 529, 1026
- Shibata, K. & Magara, T. 2011, *Living Reviews in Solar Physics*, 8, 6
- Shibata, K., Isobe, H., & Hillier, A., *et al.* 2013, *PASJ*, 65, 49
- Shibayama, T., Maehara, H., & Notsu, S., *et al.* 2013, *ApJS*, 209, 5
- Shimizu, T. 1995, *PASJ*, 47, 251
- Smith, J. C., Stumpe, M. C., & Van Cleve, J. E., *et al.* 2012, *PASP*, 124, 1000
- Strassmeier, K. G. 2009, *A&AR*, 17, 251
- Stumpe, M. C., Smith, J. C., & Van Cleve, J. E., *et al.* 2012, *PASP*, 124, 985
- Veronig, A., Temmer, M., & Hanslmeier, A., *et al.* 2002, *A&A* 382, 1070